Segmentation of Storm Enhanced Density (SED) by Boundary Flows Associated with Westward Drifting Partial Ring current

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Key Points:

- Boundary flows between Region-1 and Region-2 FACs segment SED plume into a patch.
- Localized plasma loss due to enhanced frictional heating within boundary flows.
- No external IMF direction change needed in this segmentation scenario.

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Abstract

The segmentation mechanism of polar cap patches is agreed to be related to temporal changes of IMF or transient reconnection. In this letter, using GITM driven by two-way coupled Block-Adaptive-Tree-Solarwind-Roe-Upwind-Scheme (BATSRUS) and Rice Convection Model (RCM), a new segmentation mechanism is proposed. This mechanism works as follows: a strong boundary flow between the Region 1 and Region 2 Field-Aligned Currents (FACs) develops while a shielding process develops in the inner magnetosphere. As the partial ring current drifts westward, the peak of the boundary flow also moves westward. This strong boundary flow raises the ion temperature through enhanced frictional heating, enhances the chemical recombination reaction rate and reduces the electron density. When this boundary flow crosses the Storm-Enhanced Denstiy (SED) plume, the plume will be segmented into patches. No external IMF variations or transient reconnections are required in this mechanism.

1 Introduction

Polar cap patches are high F-region plasma density structures in the polar cap where the density is much higher than the background level. Both Storm Enhanced Density (SED) and patches have drawn much attention since their discoveries, because the large density gradient at their edges is favorable for ionospheric irregularities to develop. Thus they are the major space weather concern in the polar cap that can disrupt navigation and communication signals during times of geomagnetic disturbance.

Various formation mechanisms of patch segmentation have been proposed and it is still controversial which mechanism is the dominant one. However, it is typically believed that the formation of patch needs a reservoir of high-density plasma and a segmentation mechanism. There are two sources of high-density plasma: The first source is the plasma density increase in the cusp due to soft precipitation in the F region (Goodwin et al., 2015; Rodger, Pinnock, Dudeney, Baker, & Greenwald, 1994; Walker, Moen, Kersley, & Lorentzen, 1999). The second source is the dense thermal plasma from the sunlit region (Anderson, Buchau, & Heelis, 1988). SED is a ridge of enhanced electron density in the mid-latitude region that forms during storm time (Foster, 1993), and is considered to be an important plasma source for patches. Ren, Zou, Gillies, Donovan, and Varney (2018) show that the altitude profiles of patch density and temperature resemble those of SED. That is, the patch density is usually higher than the sector median above 200 km and lower than that below 200 km, and the patch electron temperature is lower than the surrounding region. This is strong evidence that patches originate from SED.

The dense plasma of SED can be transported into the high-latitude region to form SED plume or Tongue of Ionization (TOI) (Foster, 1984). To isolate the plume into patches, a segmentation mechanism is needed. It is generally believed that the segmentation is closely related with IMF variations. Patches can be segmented due to the temporal changes in both IMF Bz (Anderson et al., 1988) and By (Milan, Lester, & Yeoman, 2002; Sojka et al., 1993). Patches can also be generated by Flux Transfer Events (FTE) near the cusp region (Carlson et al., 2006; Carlson, Oksavik, Moen, & Pedersen, 2004; Lockwood & Carlson, 1992). FTE can lead to meso-scale (about 100 km) fast flow channels near the cusp (Oksavik et al., 2006; Pinnock et al., 1993) and then bring dense plasma into the polar cap region (Lockwood & Carlson, 1992; Rodger et al., 1994; Zhang et al., 2011). Moen et al. (2006) further show that pulsed return flow due to transient reconnection can chop the plume into patches in the subauroral region. Another possibility is that fast flow jets due to transient reconnection produce high-temperature and low-density plasma through frictional heating and erode the pre-existing plume into patches (Ogawa, Buchert, Nishitani, Sato, & Lester, 2001; Valladares, Basu, Buchau, & Friis?Christensen, 1994; Valladares, Decker, Sheehan, & Anderson, 1996). Zhang et al. (2013) further show that

the low-density plasma due to precipitation can segment the TOI into patches during transient reconnection. Zhang et al. (2016) show that Subauroral Polarization Streams (SAPS) can transport low-density plasma from the night side to cut the SED during a substorm.

In this letter, Global Ionosphere Thermosphere Model (GITM) was used to simulate the geomagnetic storm on Sep 7, 2017 and study the segmentation of SED plume into a polar cap patch. GITM is driven by high-resolution coupled Block-Adaptive-Tree-Solar Wind-Roe-Upwind-Scheme (BATSRUS) and Rice Convection model (RCM) within the Space Weather Modeling Framework (SWMF) (Tóth et al., 2005; Tóth et al., 2012). Our results are not only consistent with observations, but also for the first time show that a strong boundary flow can form between the Region 1 and Region 2 FACs during the shielding development and westward drift of partial ring current in the inner magnetosphere, and this flow can reduce the electron density and segment SED plume into a patch.

2 Observations and Modeling

2.1 Model Description

GITM is a three-dimensional spherical grid that models the Earth's global ionosphere and thermosphere system self-consistently (Ridley, Deng, & Tóth, 2006). In this study, the spatial resolution of GITM is set to 1 degree in latitude and 2 degree in longitude with 50 vertical levels (from $\sim$ 100 to $\sim$ 700 km). GITM allows different models of high-latitude electric fields, auroral particle precipitation, and solar radiation as inputs. In this study, the Flare Irradiance Spectral Model (FISM) (Chamberlin, Woods, & Eparvier, 2007, 2008) is used to provide solar irradiance input for the whole simulation. FISM is an empirical model of the solar irradiance spectrum covering from 0.1 to 190 nm with 1 nm resolution and 1-minute temporal resolution. Using the high-resolution FISM model, the solar flare impact on the ionosphere and thermosphere can be captured.

Various models of high-latitude electric fields and auroral particle precipitation can be used in GITM. At first, GITM was run for two quiet days (Sep 4-5) to achieve steady state. During this stage, the Fuller-Rowell and Evans empirical model (Fuller-Rowell & Evans, 1987) and Weimer empirical electric field model (Weimer, 2005) are used as the precipitation and electric field input, respectively. The Fuller-Rowell and Evans model is driven by Hemispheric Power (HP), which is found to be linearly related with $AL^{\frac{1}{2}}$ in Østgaard, Vondrak, Gjerloev, and Germany (2002) empirical model. We use the observed AL index from the Kyoto WDC to calculate the hemispheric power. The Weimer model is driven by the observed solar wind data from OMNIWeb. Starting from Sep 6 00 UT, the high-latitude driver is switched to the electric field and precipitation from the 2-way coupled BATSRUS and RCM models, as part of the University of Michigan Space Weather Modeling Framework (SWMF) (De Zeeuw et al., 2004; Tóth et al., 2012). The coupled model is driven by the real solar wind data from OMNIWeb. High resolution grid ($1/8R_E$) in the inner magnetosphere, magnetopause and current sheet has been implemented to better capture dynamic FACs and precipitation structures.

2.2 Solar Wind and IMF

An intense double-dip geomagnetic storm occurred at Sep 7-8, 2017 and has received much attention (e.g. Aa et al., 2018; Jin et al., 2018). The SYM-H index reached -142 nT during the first dip and then $\sim$-100 nT during the second dip. The double-dip storm was driven by multiple interplanetary coronal mass ejections (ICMEs) passing the Earth successively (Shen, Luo, Chi, Xu, & Wang, 2018). From top to bottom, Figure 1a-1f shows the Interplanetary Magnetic Field (IMF) By, Bz in GSM coordinates, solar wind speed, proton number density, dynamic pressure and the SYM-H index during Sep 7-8, 2017.

The IMF Bz turned southward at $\sim$ 2040 UT on Sep 7 due to the arrival of the first ICME, leading to the initial decrease of SYM-H. At $\sim$ 2330 UT, a shock-ICME complex structure arrived with the shock enhanced IMF Bz reaching ~ -30 nT. A sudden storm commencement signaled by the sudden SYM-H increase was also initiated by this shock. After that, the SYM-H continued to decrease due to the large southward IMF and reached the first minimum at $\sim$ 0100 UT. Then the storm began to recover until the arrival of another ICME at $\sim$ 1105 UT on Sep 8, which led to the second main phase. In this study, we focus on the development and structuring of SED during the first main phase. In Figure 1f, the dashed line represents the simulated Dst index. Although there is a $\sim$ 20 nT negative bias, the temporal variation of the simulated Dst is consistent with the SYM-H index.

2.3 Model Results

The SED formation and patch segmentation processes are shown in Figure 2, and the color-coded vertical dashed lines in Figure 1 indicate the timing of the three snapshots (2100 UT, 2130 UT and 2300 UT). The first row shows the GPS TEC in geomagnetic coordinates. It is shown that the SED and the plume began to develop after the southward turning of IMF, and the segmentation of the plume occurred at 2300 UT. The bottom three rows in Figure 2 show the dynamical evolution of the ionosphere and inner magnetosphere from SWMF after the first southward turning. The second row shows the GITM electron density at 349.4 km with color contours and electric equipotential line contours every 10^4 V in geomagnetic coordinates centered at the Magnetic North Pole. The third row shows the SWMF FACs with color contours and the same electric equipotential contours every 10^4 V in geomagnetic coordinates. The green dashed line represents the open-closed field line boundary (OCB). The bottom row shows the plasma pressure in the XY plane at Z=0 in GSM coordinates. Movies of GPS TEC, electron density at 350 km, FACs and plasma thermal pressure are provided in the supporting information. Based on the comparison between the first and second row in Figure 2, it is shown that the simulation results successfully captured the development of the SED and plume after 2100 UT, and the segmentation of plume at 2300 UT.

As shown in Figure 1, the IMF Bz turned southward at 2040 UT on Sep 7, reaching -10 nT at 2100 UT, and stayed steadily southward for about 2.5 hours until the shock arrival at 2330 UT. Due to this southward turning, a two-cell convection pattern started to form and the convection electric field penetrated to mid-latitudes. The convection also tilted towards the dawn side, which is likely due to the long lasting positive IMF By (Weimer, 2005). In the equatorial plane of the inner magnetosphere, a weak nearly symmetric ring current can be seen at 21 UT. Both the Region 1 and Region 2 FACs existed. However, the Region 2 FACs were weaker than the Region 1 FACs, because of weak partial ring current. The penetrated electric field was able to reach the plasma in the sunlit mid-latitude region. Both the two-cell convection pattern and OCB expanded to lower latitude. More plasma in the mid-latitude followed the convection to higher latitude. Because the magnetic field is not strictly vertical in subauroral and auroral region, the convection flow had an upward component and transported plasma to higher altitude, where the charge exchange and recombination rate is lower (Heelis, Sojka, David, & Schunk, 2009; Zou et al., 2014, 2013). This lifting and horizontal advection led to the transport of increased electron density.

At 2130 UT, more plasma was transported from the plasma sheet to the inner magnetosphere due to the steady southward IMF, and the partial ring current developed on the night side. Because of the increase of the convection electric field, the plasma in the dayside ring current was not in a closed drift path anymore and thus drifted sunward to the magnetopause, which led to the disappearance of ring current on the day side. Together with continued buildup of ring current on the night side, the partial ring current formed with larger azimuthal pressure gradient and thus stronger Region 2 FACs accord-

ing to the Vasyliunas equation (Vasyliunas, 1970). This led to stronger Region 2 FACs which flowed up from the dawnside ionosphere and down into the duskside. The partial ring current charged up positive on the duskside and negative in the dawnside to support the Region 2 FAC (Wolf, Spiro, Sazykin, & Toffoletto, 2007). This charge separation generated a dawnward electric field, which was opposite to the convection electric field. This is referred to as the shielding phenomenon. During this period, the shielding effect began to grow. The electron density in the mid-latitude SED base region can continue to grow as long as the convection electric field there is not completely shielded, and a local maximum of electron density is created (see the electron density at 2130 UT in Figure 2). The high-density plasma then drifted further northward and into the polar cap region.

At 2300 UT, the IMF Bz had been steadily southward for two hours. A strong partial ring current built up due to the persistent and strong transport from the plasma sheet to the inner magnetosphere. The peak pressure reached ∼20 nPa and the pressure gradient generated much stronger Region 2 FACs. Due to the shielding effect, the electric potential was gradually confined to the poleward part of the Region 2 FACs, and the penetrating electric field equatorward of the Region-2 FACs weakened. The direct impact of this electric field pattern change is the formation of intense plasma flows at the boundary between the Region-1 and Region-2 FACs, commonly referred to as the Birkeland Current Boundary Flows (BCBF) (Archer et al., 2017; Archer & Knudsen, 2018). Associated with the boundary flows are increased electric field/velocity gradients on both sides. The BCBF flow may transport the low density plasma from the nightside to the dayside and also gradually erode the high density SED plume through enhanced frictional heating. As a result, a patch is segmented from the SED base. It is subsequently transported across the OCB and into the polar cap.

3 Discussion

SED segmentations were previously found to be related with the temporal change of IMF By and Bz (e.g. Anderson et al., 1988; Milan et al., 2002; Sojka et al., 1993) or transient reconnection (Carlson et al., 2006, 2004; Lockwood & Carlson, 1992). In this case, between 2100 UT and 2300 UT, both IMF By and Bz were quite steady, ruling out the possibility of temporal changes of the IMF.

To gain further insight into the magnetospheric origin of convection flows cutting the patch, a 3-D view of the simulated magnetospheric configuration at 2300 UT when the SED is segmented is shown in Figure 3. Color contours of plasma thermal pressure are shown in the cut plane XY at Z=-1 RE. Colors on the sphere at 3 RE are contours of the FACs (positive values mean upward FACs). Three white lines represent the sample field lines at the center of the SED base (MLAT=57.3, MLT=15.0), center of the boundary flows (MLAT=63.0, MLT=13.8), and patch (MLAT=63.5, MLT=13.4). The center of the SED base is where the electron density at 350 km reaches the local maximum, the center of the boundary flow is where the electron density reaches the local minimum, and similarly the center of patch is where the density reaches local maximum. The footprints of these magnetic field lines are marked as white stars in Figure 2. The polar cap patch was on open magnetic field lines and connected to the solar wind, showing that its poleward movement was due to the magnetospheric convection driven by the southward IMF. The SED base was mapped to the inner edge of the partial ring current, where the undershielded electric field was responsible for the SED lifting and TEC growth. This also explains why the SED base stopped providing plasma to the auroral region after the shielding further developed. The boundary flow region was mapped to the outer boundary of the partial ring current (poleward boundary of Region 2 FACs in the ionosphere). This proves that the segmentation was due to the boundary flows located between the Region 1 and Region 2 FACs.

We further investigate how the boundary flows segment the SED plume into a patch. The plasma parcel at the center of the boundary flow at 2300 UT was traced backward in time to 2130 UT to identify its source region. The center of the boundary flow is located at MLAT=63.0,MLT=13.8 and marked with a white star in Figure 4c. It also co-located with the minimum of the electron density. The altitude profile of the electron density at 2300 UT is represented by the blue curve in Figure 4d. When the plasma parcel was traced back in time, it was found that the source region was located at MLAT=58.3, MLT=15.6 at 2130 UT and also marked with a star in Figure 4a. The plasma parcel experienced two development phases, i.e., a growing phase and a decaying phase. The growing phase started at 2130 UT and continued to 2225 UT. The corresponding altitude profiles of the electron density at these two time cadences are represented by the red and black curves in Figure 4d, respectively. During the growing phase, the plasma was lifted to higher altitudes, i.e., the hmF2 increased from 270 km to 320 km. The lifting was found to be mainly due to projection of the northward convection flows in the vertical direction. At the end of the growing phase, this plasma parcel was located at MLAT=60.1,MLT=15.0, indicated by the star near the equatorward boundary of the boundary flow shown in the central panel of Figure 4b. The electron density at the hmF2 ($\sim$ 300 km) reached its maximum of $4.09\times10^{11}m^{-3}$ at 2225 UT. The decaying phase started at 2225 UT and the electron density in the F-layer below 400 km region kept decreasing. The temporal evolution of the horizontal ion drift, electron density and ion temperature in the plasma parcel at around 350 km is shown in Figure 4e-g, respectively. The horizontal flow was mainly westward and increased during the whole period due to the constant southward IMF, continuously growing particle ring current and the associated Region-2 current system. The electron density reached the maximum at 2225 UT. After 2225 UT, the convection flow continued to increase, and the ion temperature also increased due to stronger frictional heating between the ions and neutrals, which also resulted in an electron density decrease.

The loss mechanism is further quantified. When the plasma parcel moved from the source region to the boundary flow region, at 300 km (around hmF2) the velocity difference between ion and neutral increased from $\sim$100 m/s to $\sim$600 m/s, the ion temperature increased from $\sim$1400 K to $\sim$1700 K, and the neutral temperature remained at $\sim$1400 K. Based on the temperature dependent dissociative recombination equations in St.-Maurice and Torr (1978), the charge exchange rate between O^+ and N_2 increases from $5.73 \times 10^{-13}cm^3s^{-1}$ to $9.91\times10^{-13}cm^3s^{-1}$, when the ion temperature increases from $\sim$1400 K to $\sim$1700 K. The N_2 density was $\sim 1.5\times10^{14}m^{-3}$ and O^+ density was $\sim 3\times10^{11}m^{-3}$ in the model output. The production rate of NO^+ thus increased $1.88 \times 10^7m^{-3}s^{-1}$. The total NO^+ conversion during the decaying phase ($\sim$ 35 min) is $4\times10^{10}m^{-3}$. Because the recombination between NO^+ and e^- is rapid, the electron density loss should also be $4\times10^{10}m^{-3}$, which matches the amount of electron lost between 2225 UT and 2300 UT in the model output. This simple calculation confirms that the enhanced chemical reaction within the enhanced boundary flows led to the electron density decrease and the segmentation of the SED plume into a patch.

There are several proposed patch segmentation mechanisms with convincing evidence reported in previous literature, which highlights the complicated nature of the high-latitude ionospheric electrodynamics. In this numerical simulation, the partial ring current formed has a single pressure peak and thus the associated FAC and flow pattern are relatively simple. However, the ring current can have fine structures (Buzulukova et al., 2010; Wei, Yu, & He, 2019) and the associated FAC and flow pattern can be more complicated, which can potentially lead to more structures in the segmented patch. A local maximum of the convection flow and a local density minimum at 2235 UT in Figure 4e-g further confirms this idea.

4 Summary and Conclusions

A new segmentation mechanism of SED plume into polar cap patches is proposed through numerical model study. A simulation of the ionospheric response to the geomagnetic storm on Sep 7, 2017 has been carried out using the Global Ionosphere and Thermosphere Model (GITM) driven by the two-way coupled BATSRUS and RCM within SWMF. The simulation results are validated by comparison to GPS TEC observations and SYM-H index.

During storm time, dense SED plasma from the mid-latitude ionosphere is transported by the expanded two-cell convection pattern into the high-latitude region to form the SED plume. Meanwhile, plasma from the plasma sheet is transported into the inner magnetosphere by strong earthward convection and forms a strong partial ring current centered on the nightside. The partial ring current closes through the Region 2 FACs and gradually shields the subauroral region from the high-latitude penetrating electric fields. As the partial ring current grows, the shielding effect strengthens, causing the boundary flows between the Region 1 and Region 2 FACs to increase. This leads to enhanced ion temperature and chemical recombination rate, and thus reduced electron density. When the center of the partial ring current drifts westward due to the gradient and curvature drifts, these large boundary flows also move westward from the night side to the day side. When they encounter the SED plume, the plume is segmented into a patch, which later move further into the polar cap. During this process, no IMF variations or transient reconnections are required. The segmentation mechanism is entirely due to magnetospheric internal processes.

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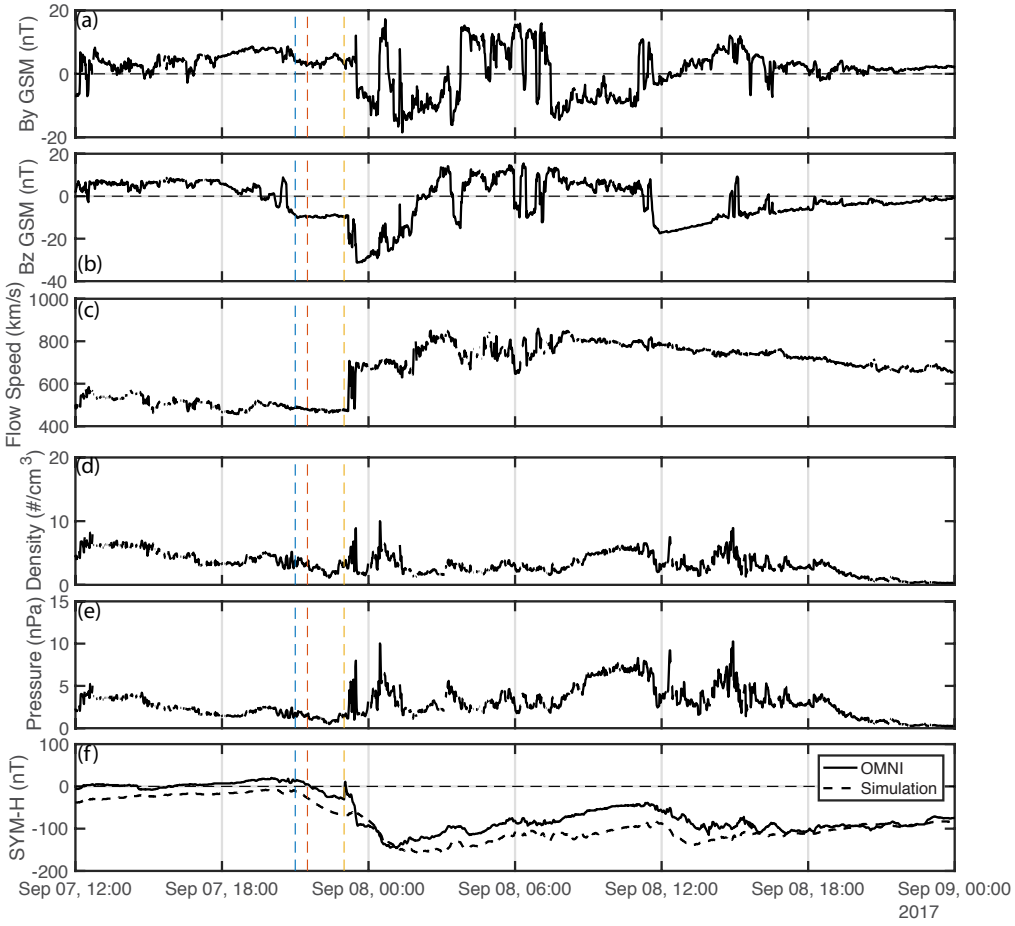

Figure 1. Solar wind data and geomagnetic index from 12 UT Sep 7 to 0 UT Sep 9, 2017. (a) IMF By component in the GSM coordinates, (b) IMF Bz component in the GSM coordinates, (c) solar wind speed, (d) proton number density, (e) solar wind dynamic pressure, and (f) the SYM-H index (solid line) and simulated Dst index (dash line). The color-coded vertical dashed lines indicate timing of the three snapshots in Figure 2.

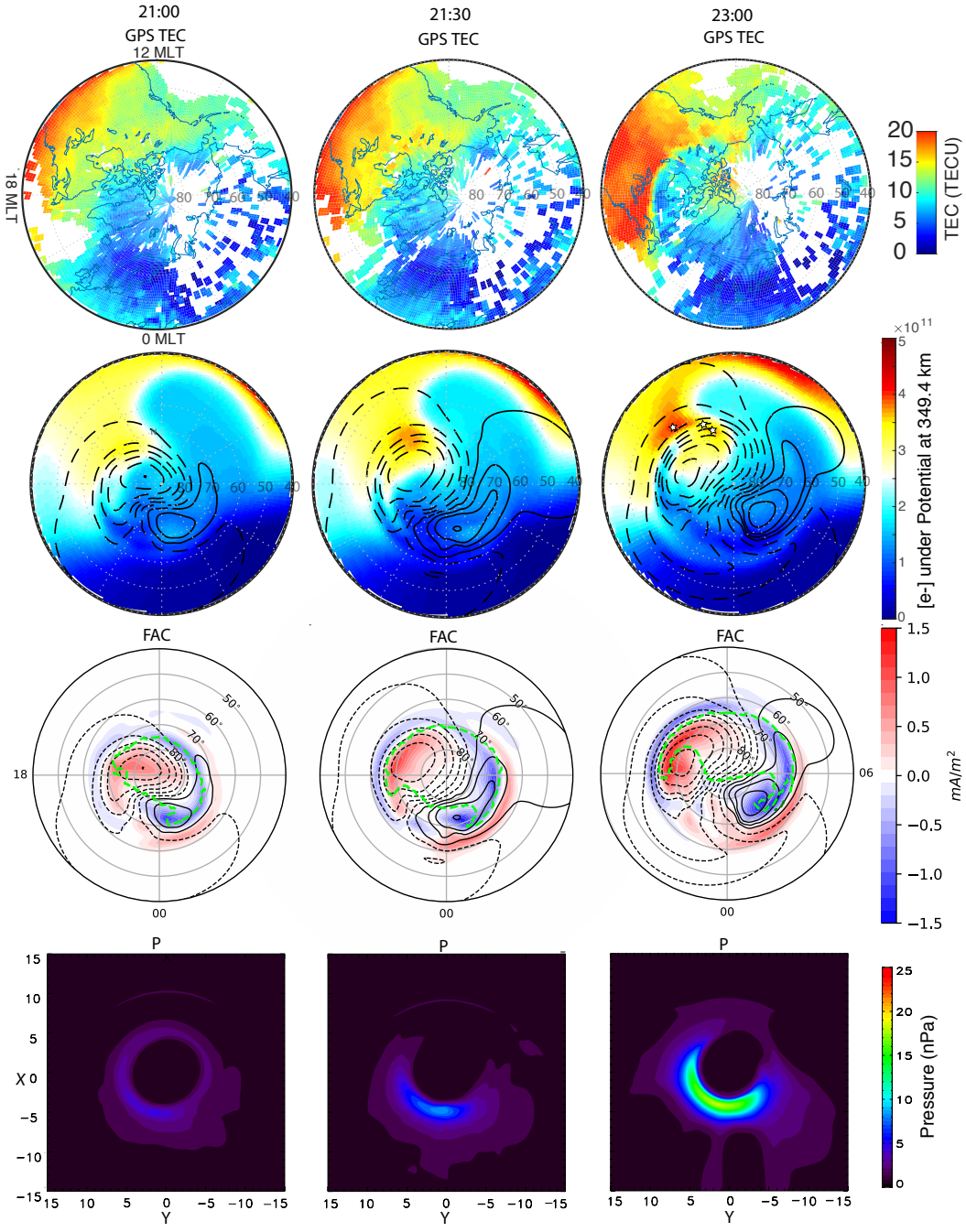

Figure 2. First row: Polar view of TEC in the geomagnetic coordinates. Second row: Polar view of modeled electron density (color contours) at 349.4 km under electric equipotential (contour lines) in the geomagnetic coordinates. The SED base, boundary flow and polar cap patch are marked with stars at 23 UT. Third row: Polar view of the modeled FACs (color contours) under electric equipotential (contour lines) in the geomagnetic coordinates. The green dashed line represents the open-closed field line boundary. Fourth row: the modeled plasma thermal pressure in the z=0 plane in the GSM coordinates.

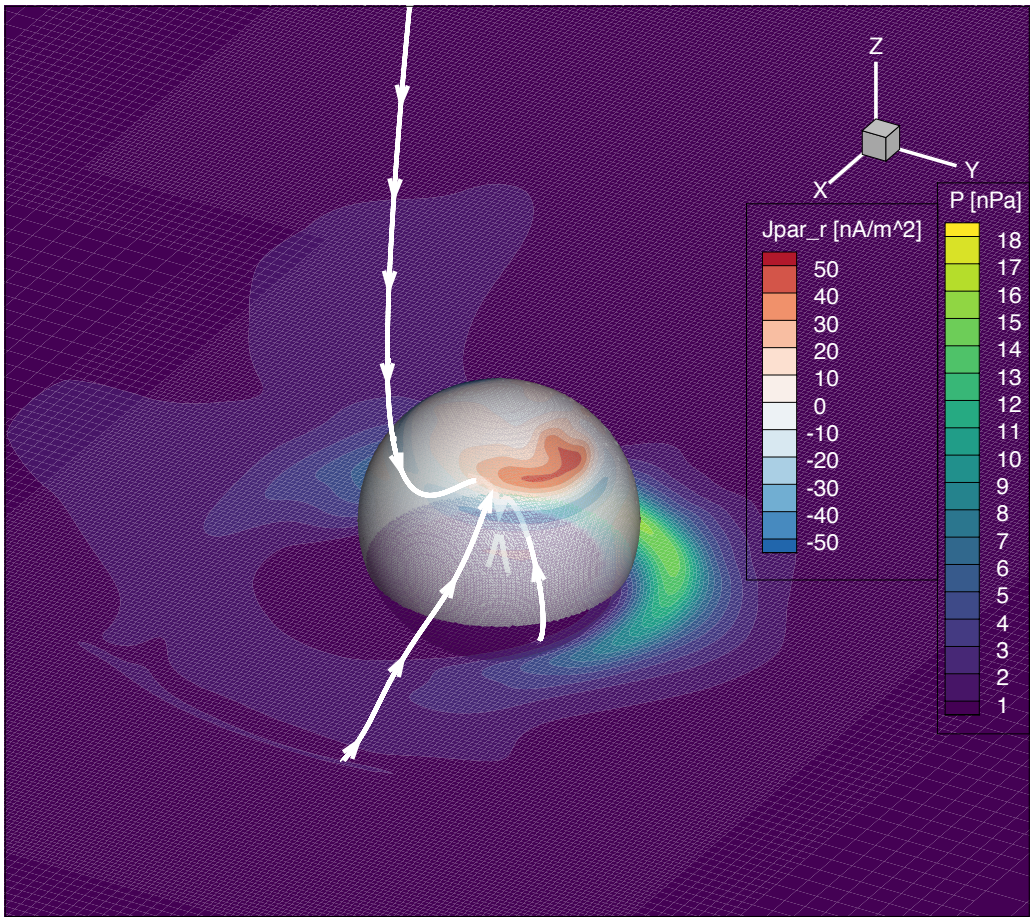

Figure 3. A 3-D view of the simulated magnetosphere extracted from the time step when SED was segmented at 23 UT. Shown in the cut plane XY at Z= -1 RE are color contours of the plasma thermal pressure. Color contours on the sphere with a radius of 3 RE indicate the field-aligned current density (positive values mean upward field-aligned currents). Three white lines represent the sample magnetic field lines at the SED base (the lowest lat.), boundary flow (the middle lat.) and patch (the highest lat.) around this time. The foot prints of these magnetic field lines are shown as stars in Figure 2.

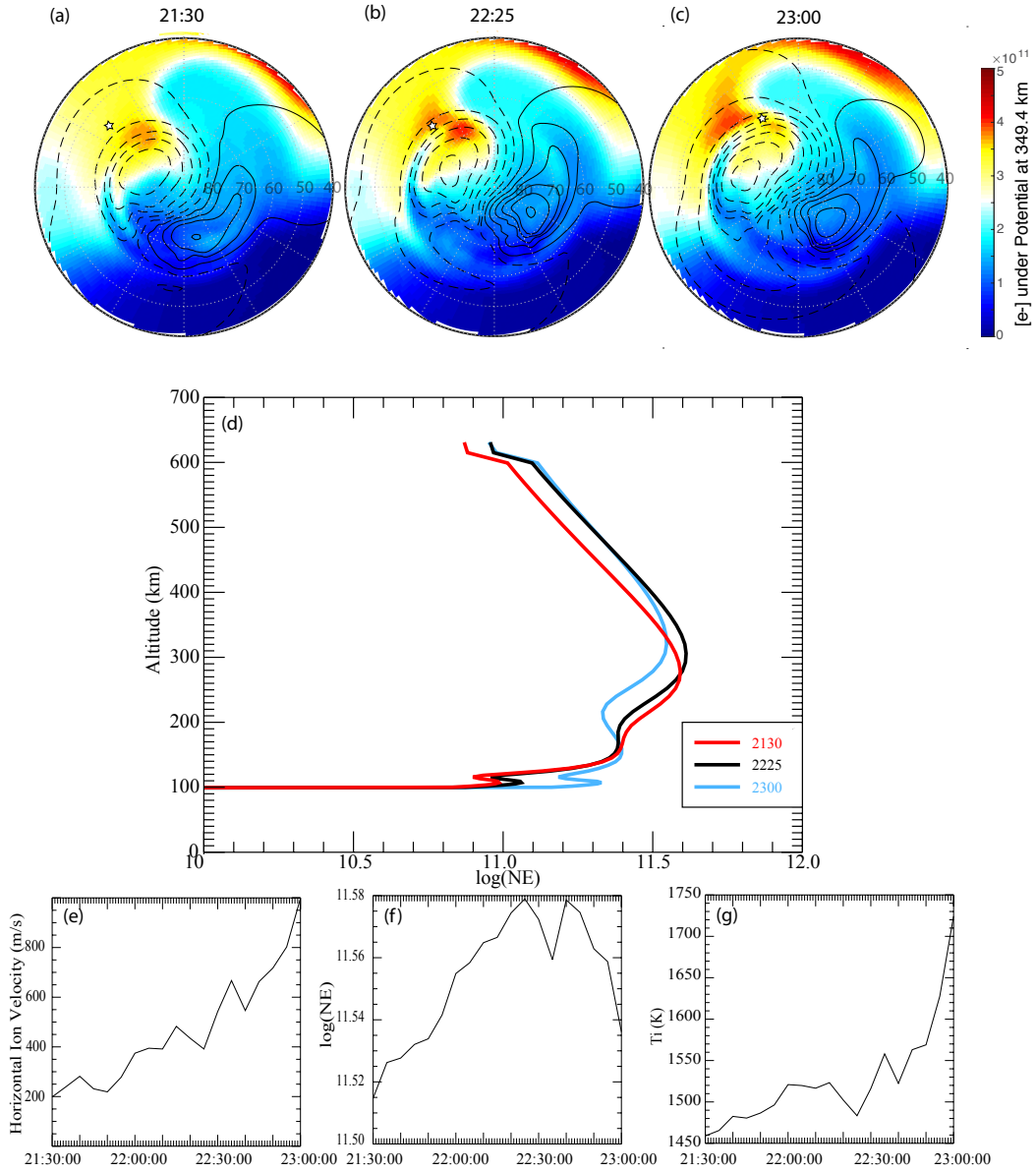

Figure 4. One simulation plasma parcel is traced backward in time. The locations of the plasma parcel and altitude profiles of the electron density at three different time cadences are shown. (a)-(c) Polar view of the F-region electron density and the locations of the traced plasma parcel are marked with stars at 2130, 2225 and 2300 UT, respectively. The format of polar view plots of the electron density is the same as the second row in Figure 2. (d) The red curve indicates the initial density profile at 2130 UT, the black curve indicates the density profile at the end of the growing phase, i.e., 2225 UT, and the blue curve indicates the density profile at the end of the decaying phase, i.e., 2300 UT. (e)-(g) The temporal evolution of the horizontal ion drift, electron density and ion temperature in the plasma parcel at around 300 km.

Figure 1.

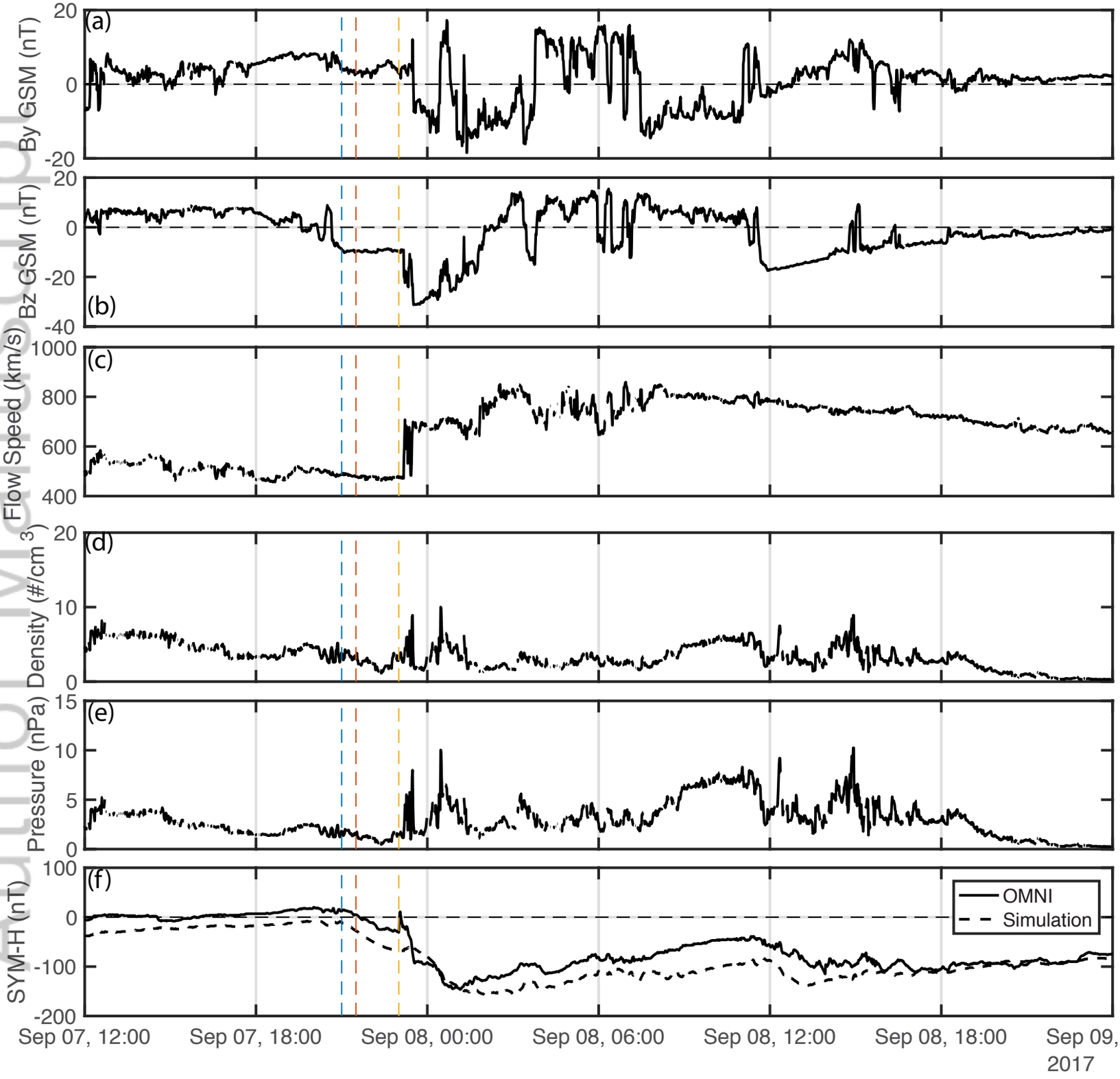

Figure 2.

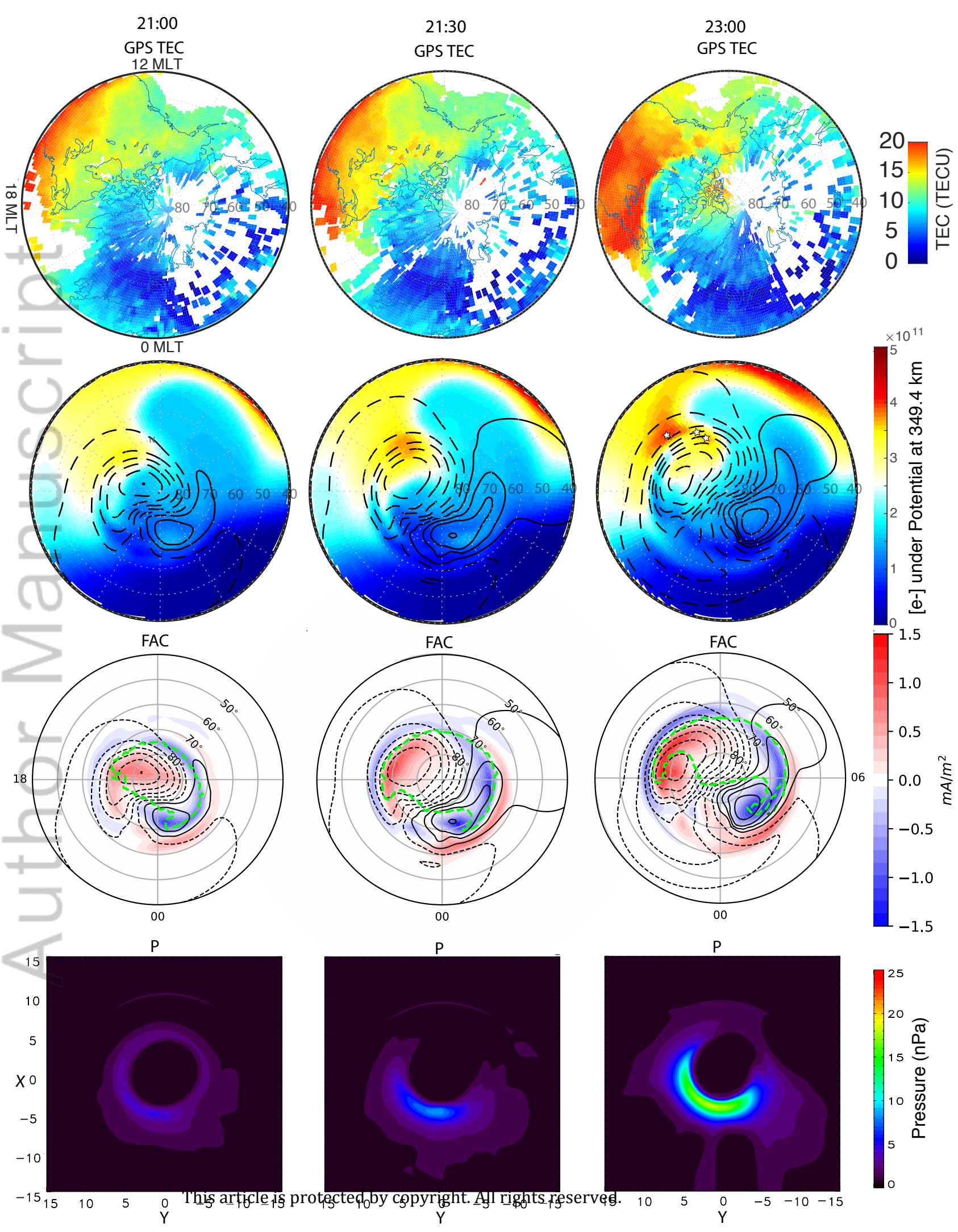

Figure 3.

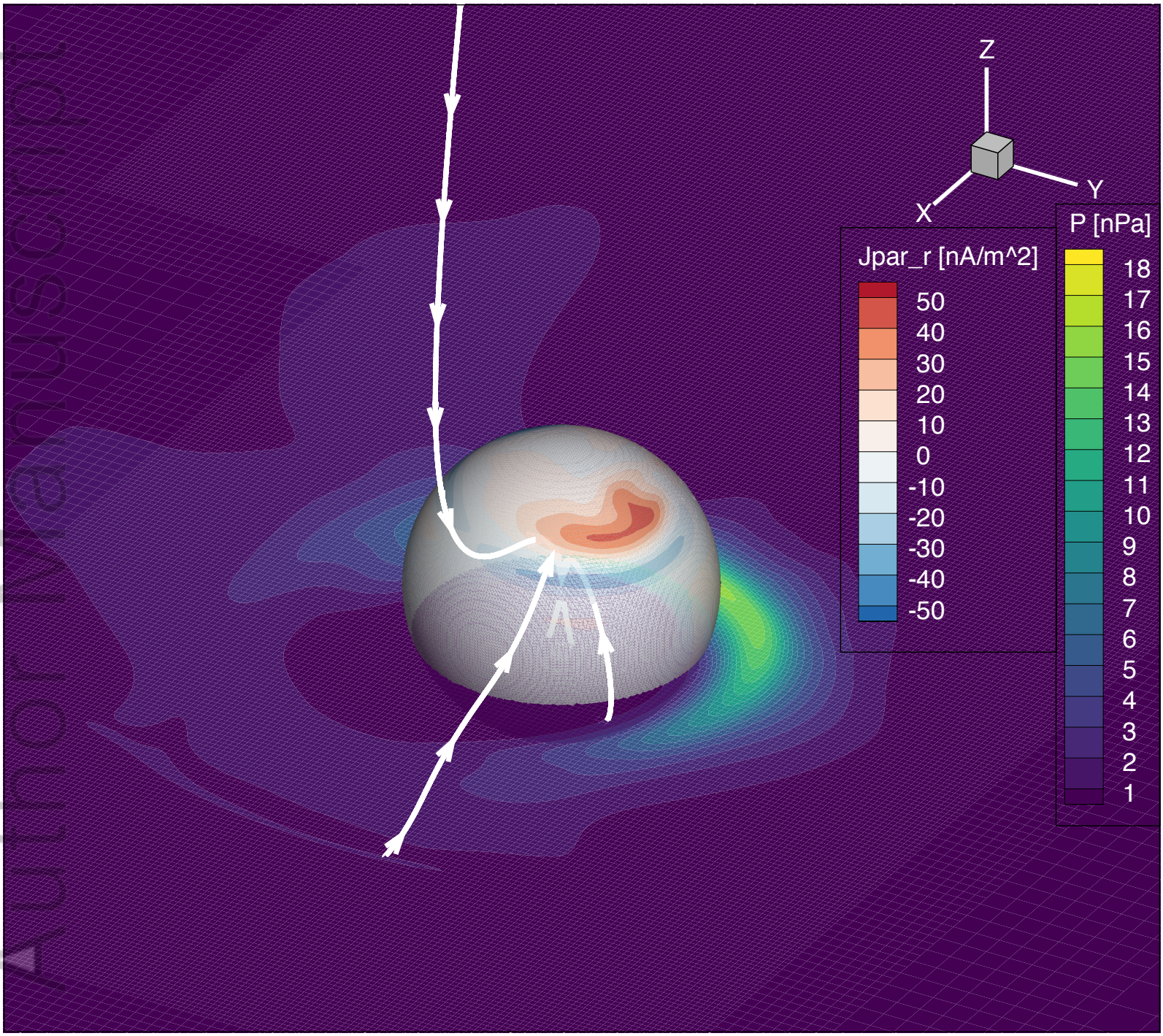

Figure 4.