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### Key Points:

- Modulation of whistler mode waves by ion-scale waves are observed in the distant magnetotail
- The whistler mode waves are excited by field-aligned electron populations generated by the ion-scale waves
- The ion-scale waves, possibly kinetic Alfvén waves, are driven by tailward streaming ion beams

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## Modulation of Whistler Mode Waves by Ion-Scale Waves Observed in the Distant Magnetotail

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**Abstract** Wave activities in tailward flows have been explored in the distant magnetotail at  $\sim 54 R_E$ , where the coupling between whistler mode waves and ion-scale waves was observed. The whistler mode waves periodically appeared at each cycle of the ion-scale waves, and the electron distribution functions associated with the whistler mode waves showed enhancement in the direction parallel to background magnetic field. Wave analyses show that the field-aligned electron components act as the energy source of the whistler mode waves. The ion-scale waves are generated by the interaction of the hot ion beam with background ions. A likely candidate of the ion-scale wave is the kinetic Alfvén wave, which can generate the enhanced field-aligned electron populations by the parallel electric field.

**Plain Language Summary** The ionized particles constrained by the Earth's magnetic field involve dynamics of different spatial and temporal scales. Because of large mass difference, the scales characterizing the ion and electron dynamics are very different, and the related processes are often considered separately. In this study, we have found that the electron scale waves can be modulated by ion-scale waves, indicating that the dynamics of electrons and ions are coupled. The ion-scale waves are induced by the streaming ion beams superposed on the background ions. In the process, electrons are accelerated parallel to the magnetic field direction by the ion-scale wave, resulting in the deviation of electron populations from the equilibrium state. The unstable electrons then release the excess energy by generating electron scale waves, called whistler mode waves.

## 1. Introduction

The plasma sheet (Bame et al., 1967) in the Earth's magnetotail is a region full of plasma activities characterized by high-speed plasma flows (Angelopoulos et al., 1992; Baumjohann et al., 1990), magnetic structures (Nakamura et al., 2002; Slavin et al., 2003), wave activities (Khotyaintsev et al., 2011; Zhao et al., 2016), and particle accelerations (Fu et al., 2011; Zhou et al., 2010). During the active times in the plasma sheet, both earthward and tailward flows are generated (Nagai et al., 1998). The earthward flows carry energy and magnetic flux to the inner magnetosphere and produce various space physics phenomenon, like energetic particle injection (Gabrielse et al., 2017; Sergeev et al., 2012), geomagnetic pulsations (Runov et al., 2014), and brightening of auroras (Donovan et al., 2008). The tailward flows stream down the magnetotail to distances of  $\sim 60 R_E$  and beyond (Artemyev et al., 2017; Kiehas et al., 2018), where the wave and particle behaviors show unique signatures for different plasma environment and magnetic field configuration (Li et al., 2014; Zhao et al., 2017; Zhou et al., 2015). However, the plasma dynamics in the distant magnetotail have been seldomly explored compared to the near earth regions.

During the plasma sheet activities, whistler mode waves are often observed in bursty bulk flows where the magnetic flux is compressed (Viberg et al., 2014; Zhang et al., 2018). Actually, the whistler mode wave exists throughout the magnetosphere (Li et al., 2013; Wilson et al., 2011). In the inner magnetosphere, the whistler mode waves often appear as broadband hiss (Summers et al., 2014) and chorus waves with rising and falling tone structures (Keika et al., 2012; Sazhin & Hayakawa, 1992). The pitch-angle scattering of the electrons by the whistler mode waves could precipitate the radiation belt electrons into the ionosphere (Li et al., 2019;

Mozer et al., 2018; Thorne et al., 2010) and results in the loss of electrons from the plasma sheet and carries energy away (Li et al., 2013; Zhang, Angelopoulos, et al., 2019).

Besides interaction with electrons, whistler mode waves could also be modulated by low-frequency waves since the inhomogeneity of the plasma density and magnetic field could influence the generation and propagation of this wave (Karpman & Kaufman, 1982; Streltsov et al., 2006). In the radiation belt region, whistler mode waves can be modulated by compressional ultralow-frequency waves, which induce electron density, magnetic field, and resonant electron variations (Li et al., 2011; Zhang, Chen, et al., 2019). Tenerani et al. (2012) have shown that ion-scale solitary waves can trap the whistler mode waves at the wave valley or hump, and this process can connect the electron-scale and ion-scale dynamics. In addition, Colpitts et al. (2016) have reported observations of coupling between whistler mode wave and electromagnetic ion cyclotron waves. However, the modulation mechanisms of whistler mode waves we present in this paper have not been addressed previously.

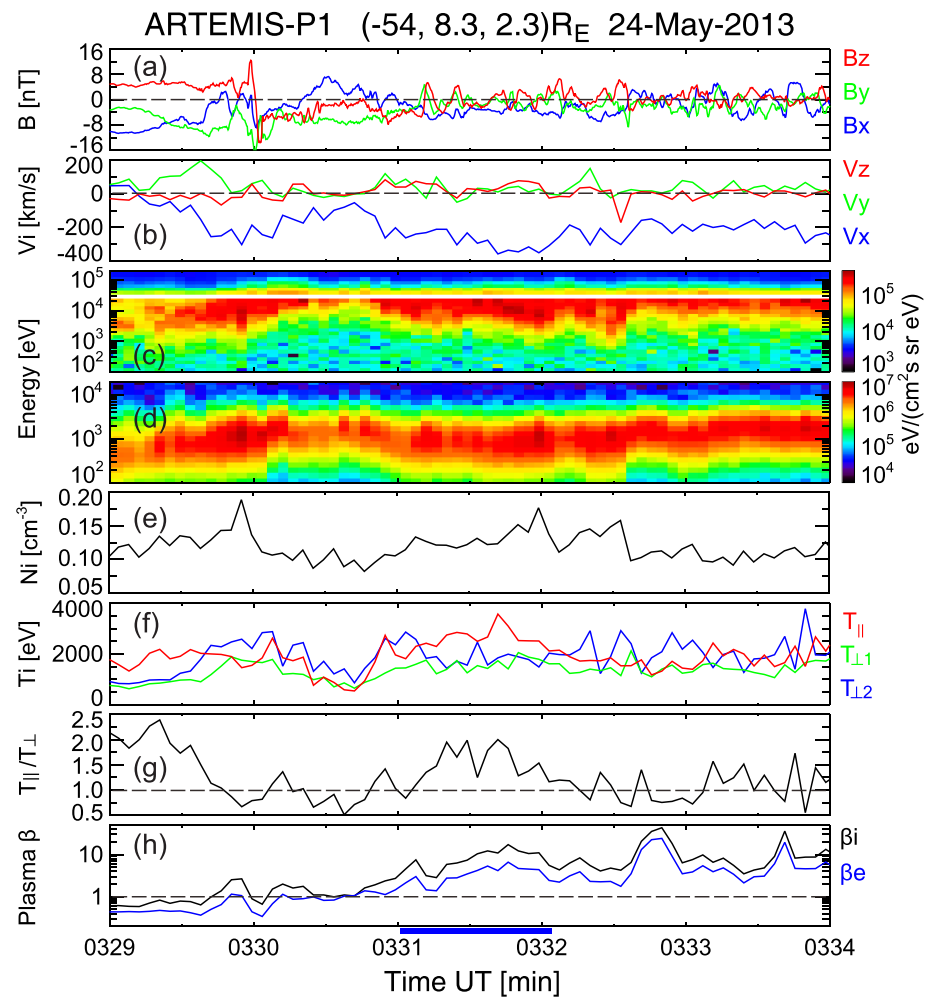
The whistler mode waves have been extensively explored in the near-Earth region (Tyler et al., 2019; Viberg et al., 2014) but not in the distant magnetotail where the plasma environment is very different. This paper presents observations of whistler mode waves at  $\sim 60 R_E$  down the geomagnetic tail. These waves are periodically excited by electron distribution functions modulated by low-frequency waves on the ion dynamic scale. To the best of our knowledge, the characteristic signatures we will show have not been reported in the magnetotail region, and they are important for enhancing our understanding of cross-scale wave interaction mechanisms in plasmas.

## 2. Observations

The Acceleration, Reconnection, Turbulence, and Electrodynamics of the Moon's Interaction with the Sun (ARTEMIS) mission (Angelopoulos, 2011) is an expansion of the Time History of Events and Macroscale Interactions during Substorms (THEMIS) mission (Angelopoulos, 2008) with two spacecraft placed at the lunar orbit, making it possible to explore the magnetotail at  $\sim 60 R_E$ . The data used in this paper come from the fluxgate magnetometer (Auster et al., 2008), electrostatic analyzer (McFadden et al., 2008), solid state telescope (Angelopoulos, 2008), and the search coil magnetometer (Roux et al., 2008). Wave spectra are calculated onboard the spacecraft by fast Fourier transforms (Cully et al., 2008). All quantities in this paper are presented in the geocentric solar magnetospheric system, unless otherwise noted.

On 24 May 2013, P1 probe of the Acceleration, Reconnection, Turbulence, and Electrodynamics of the Moon's Interaction with the Sun mission was located at  $(-54, 8.3, 2.3) R_E$  in the distant magnetotail, where a tailward plasma flow and fluctuating magnetic field were detected (Figure 1). The flow, which begins at  $\sim 032930$  UT (Figure 1b), was associated with an antidipolarization front (Li et al., 2014; Nakamura et al., 2002) near the leading edge at  $\sim 033000$  UT (Figure 1a). Inside the plasma flow, the magnetic field shows large amplitude fluctuations (Figure 1a), indicating the presence of active plasma dynamics. The ion (Figure 1c) and electron (Figure 1d) spectra in the flow are similar to that in the plasma sheet (Bame et al., 1967). From  $\sim 033100$  UT to  $\sim 033200$  UT (blue line), the ion temperature showed a much hotter component parallel to the magnetic field (Figure 1f) and resulted in the temperature anisotropy  $T_{\parallel}/T_{\perp}$  exceeding 1.0 and approaching 2.0 at some point (Figure 1g). During this interval, the ion beta  $\beta_i$ , which is defined as the ratio of ion thermal pressure to magnetic pressure, increased to  $\sim 10$  (Figure 1h). The behavior of wave activities observed during this interval is the major interest of this study.

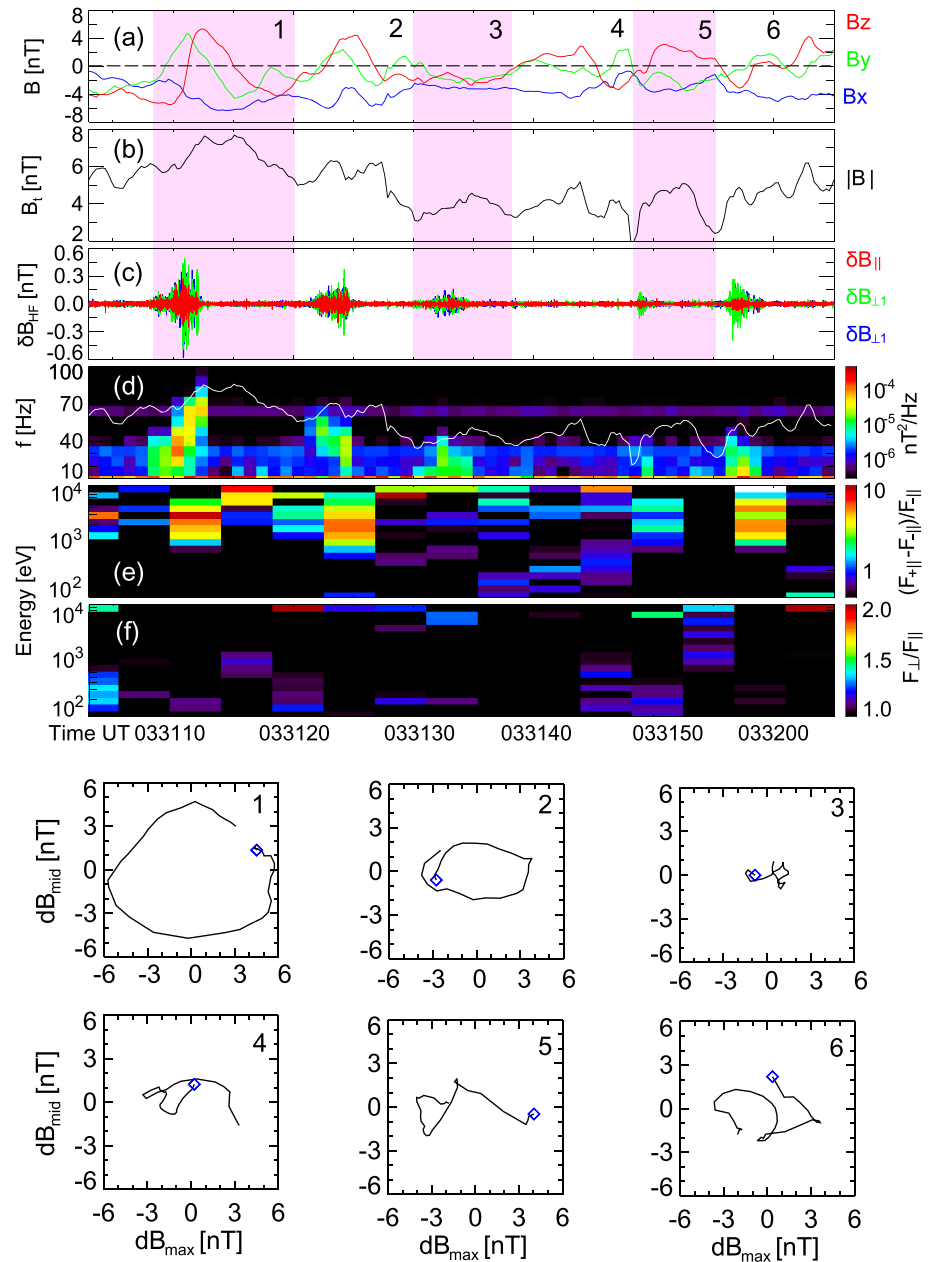
The wave activities of different frequencies were analyzed in detail from 033103 UT to 033205 UT as presented in Figure 2. Large amplitude waves of  $\sim 6$  nT have been observed from the magnetic field fluctuations (Figures 2a and 2b), and each cycle of the wave was marked as 1–6 in Figure 2a. The average frequency of this wave is about 0.1 Hz in the spacecraft frame, close to the local ion cyclotron frequency  $f_{ci} \sim 0.075$  Hz; hence, we call this wave ion-scale wave. Figure 2c shows the high-frequency magnetic field fluctuations filtered above 10 Hz, and Figure 2d shows the power spectrum of the waves calculated onboard the spacecraft with the white line indicating 0.4 times the local electron cyclotron frequency ( $f_{ce}$ ). A clear signature is that the high-frequency waves appeared periodically at each cycle of the ion-scale wave (except cycle 4), and we interpret them as modulation by the ion-scale wave.



**Figure 1.** Summary plot of ARTEMIS P1 observations in the distant magnetotail. (a) Magnetic field, (b) ion bulk velocity, (c) ion and (d) electron differential energy flux, (e) ion density, (f) ion temperature in the field-aligned coordinate system, (g) ion temperature anisotropy, and (h) ion beta ( $\beta_i$ ; black line) and electron beta ( $\beta_e$ ; blue line). The blue horizontal line below (h) marks the time interval presented in Figure 2. All the data recorded by the electrostatic analyzer instrument are under burst mode.

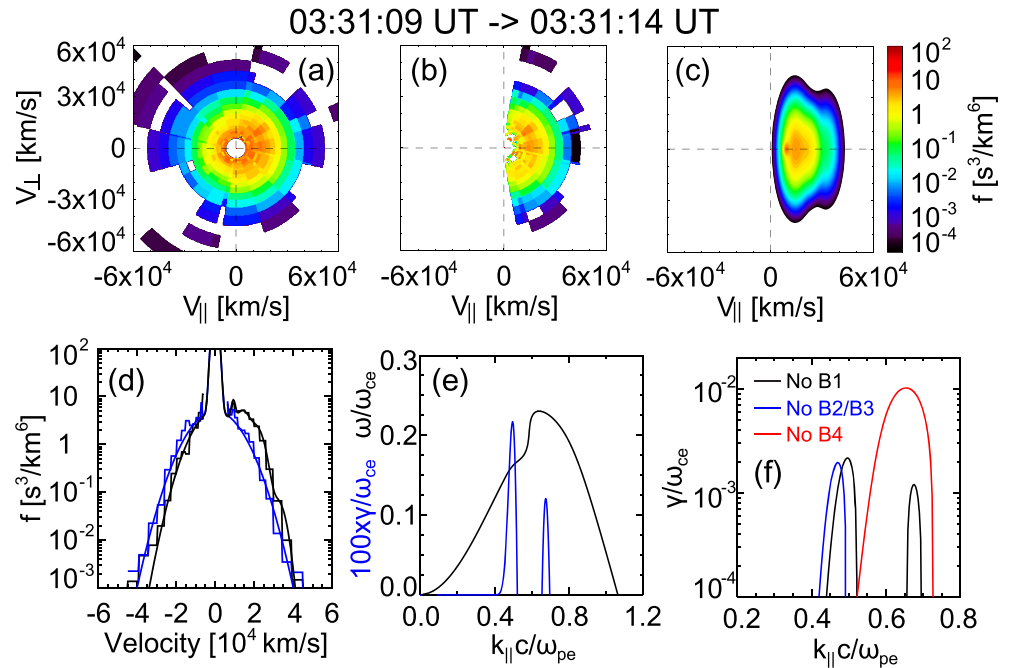
The high-frequency waves have been identified as the lower-band whistler mode waves because they are nearly right-hand circularly polarized (not shown) with a frequency lower than  $0.5 f_{ce}$  (Figure 2d). These waves propagated with angles from  $5^\circ$  to  $34^\circ$  with respect to the local magnetic field direction. The electron distribution functions associated with the whistler mode waves showed anisotropies in the direction parallel and antiparallel to the magnetic field. We measured these anisotropies by  $(F_{+||} - F_{-||})/F_{-||}$  as shown in Figure 2e, where  $F_{+||}$  ( $F_{-||}$ ) is the particle flux parallel (antiparallel) to the background magnetic field direction. The wave emissions are associated with the parallel particle flux enhancement of the electrons with energy from  $\sim 1$  to  $\sim 10$  keV (the flux anisotropy of the third whistler mode wave is weakened after normalization), while the anisotropies are not as pronounced outside the wave intervals. The transverse flux anisotropy of  $F_{\perp}/F_{||}$  ( $F_{\perp}$  is the perpendicular particle flux, and  $F_{||}$  is the average particle flux in parallel and antiparallel directions) is also calculated in Figure 2f, and there are no recognizable anisotropies when the waves appear. Thus the increased parallel electron fluxes are suggested to be the energy source of the whistler mode waves as discussed in the next section.

The hodograms of the ion-scale waves are shown in the boxes below Figure 2f, and the numbers 1–6 correspond to the six wave cycles. The first and second cycles are circularly and elliptically left-hand polarized in the spacecraft frame while the other cycles are irregularly polarized. An interesting feature to note is that the



**Figure 2.** Wave analyses and corresponding particle behaviors. (a) Magnetic field components, (b) total magnetic field strength, (c) magnetic field fluctuations filtered above 10 Hz in field-aligned coordinate system, (d) magnetic field wave spectrum with the white line that marks  $0.4 f_{ce}$ , and (e) electron flux anisotropy in parallel and antiparallel directions calculated from the parallel electron differential flux ( $F_{+||}$ ; pitch angle =  $0-30^\circ$ ) minus the antiparallel electron differential flux ( $F_{-||}$ ; pitch angle =  $150-180^\circ$ ) and then normalized by  $F_{-||}$ . (f) Electron flux anisotropy between perpendicular ( $F_{\perp}$ ; pitch angle =  $60-120^\circ$ ) and parallel ( $F_{||} = (F_{+||} + F_{-||})/2$ ) directions calculated by  $F_{\perp}/F_{||}$ . The pink and white colors separate the six wave cycles marked by 1–6 in panel (a). The boxes marked by 1–6 below panel (f) are the hodograms corresponding to 1–6 of the wave cycles. The normal direction of the hodograms points out of the paper, and the blue diamond is the starting point.

first and second ion-scale wave cycles are associated with whistler mode waves of rising and falling band structures in the wave spectrum (Figure 2d), while the whistler mode waves corresponding to other cycles are structureless. Although the magnetic field strength during the first whistler mode wave (033108 UT–033112 UT) increased from  $\sim 6$  to  $\sim 8$  nT, the wave frequency enhancement from  $\sim 30$  to  $\sim 80$  Hz is much larger and seems not to be formed by the strengthening of the magnetic field. Since the resolution



**Figure 3.** Electron distributions and wave dispersion studies of whistler mode wave 1. (a) Two-dimensional electron velocity distribution functions from 033109 UT to 033114 UT, where  $V_{\parallel}$  is the velocity parallel to the magnetic field ( $\mathbf{B}$ ) and  $V_{\perp}$  is the velocity perpendicular to  $\mathbf{B}$  and the ion bulk velocity. (b) Parallel enhanced electrons calculated by the parallel electron phase space density minus the antiparallel value; (c) fitting results of the parallel increased electrons; (d) 1-D cut (step lines) and fitting (curve line) result of the electron distributions in parallel (black line) and perpendicular (blue line) directions; (e) frequency ( $\omega$ ) and growth rate ( $\gamma$ ) as a function of parallel wave number ( $k_{\parallel}$ ) calculated based on the fitted electron distribution functions,  $\omega$  and  $\gamma$  are normalized by  $\omega_{ce}$ , and  $k_{\parallel}$  is normalized by electron plasma frequency ( $\omega_{pe}$ ) versus the speed of light ( $c$ ); and (f) growth rate calculated for the electron distribution functions without beam 1 (black line), beam 2 or beam 3 (blue line), and beam 4 (red line).

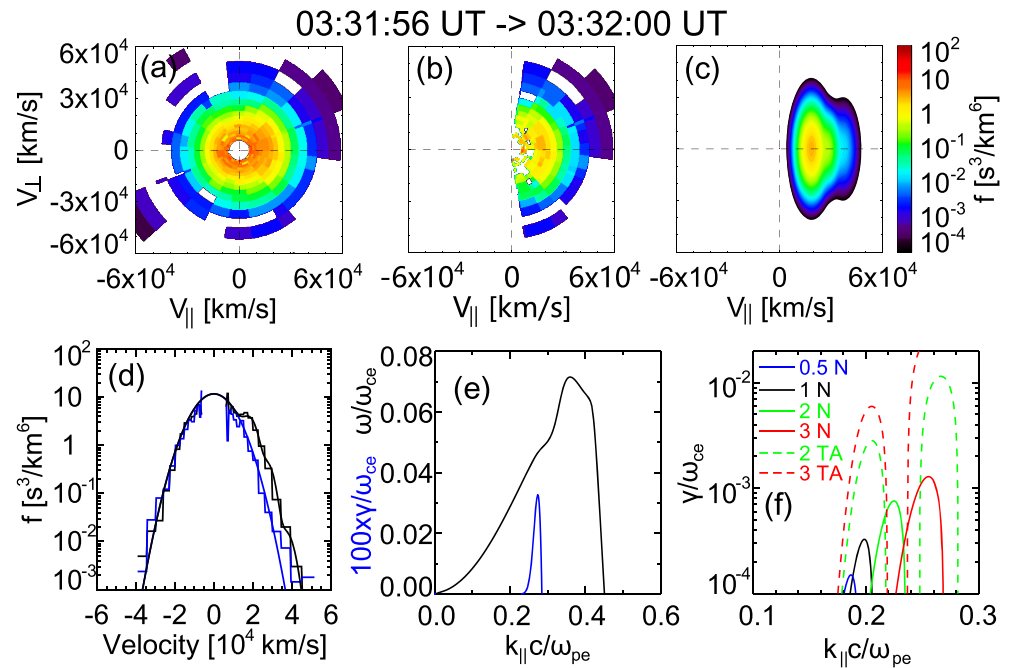
of the wave data during this interval was not sufficient, and the single probe detection could not distinguish spatial and time variations, the detail mechanisms are still awaiting further exploration. The wave features of the first two cycles indicate nonlinear coherent processes that connect the electron, and ion-scale dynamics have taken place.

### 3. Particle Distributions and Wave Analysis

To study the generation mechanism of the whistler mode waves and ion-scale waves, linear theory of hot plasma dispersion analyses based on the particle distribution functions has been employed with PDRK kinetic dispersion relation solver (Xie & Yong, 2016), and the major results are also checked with waves in homogeneous, anisotropic, multicomponent plasma (WHAMP) method (Ronmark, 1982). The electron and ion distribution functions are fitted by multi-Maxwellian-shaped distributions to make computation. When fitting the particle distributions, a cold component will be added if necessary to keep the fitted density the same as the measured value, and the drift velocity of all the populations have been shifted together to keep zero current in the system. The analysis of the first and fifth whistler mode waves are shown in Figures 3 and 4, while Figure 5 shows the results of the ion-scale waves.

The electron distribution function associated with the first whistler mode wave (Figure 3a) exhibits a higher electron phase space density in the parallel direction than in the antiparallel direction. To extract the enhanced component, the net population is calculated as the parallel phase space density minus the antiparallel value, as shown in Figure 3b. The enhanced electrons are diffused on the right side of the phase space (Figure 3b) and could not be represented by a single beam; hence, multicomponent drifting bi-Maxwellian distributions have been used to fit this population (Figure 3c). The fitting results in parallel and perpendicular directions are shown in Figure 3d, and the fitting parameters are listed in Table 1 in the first set.

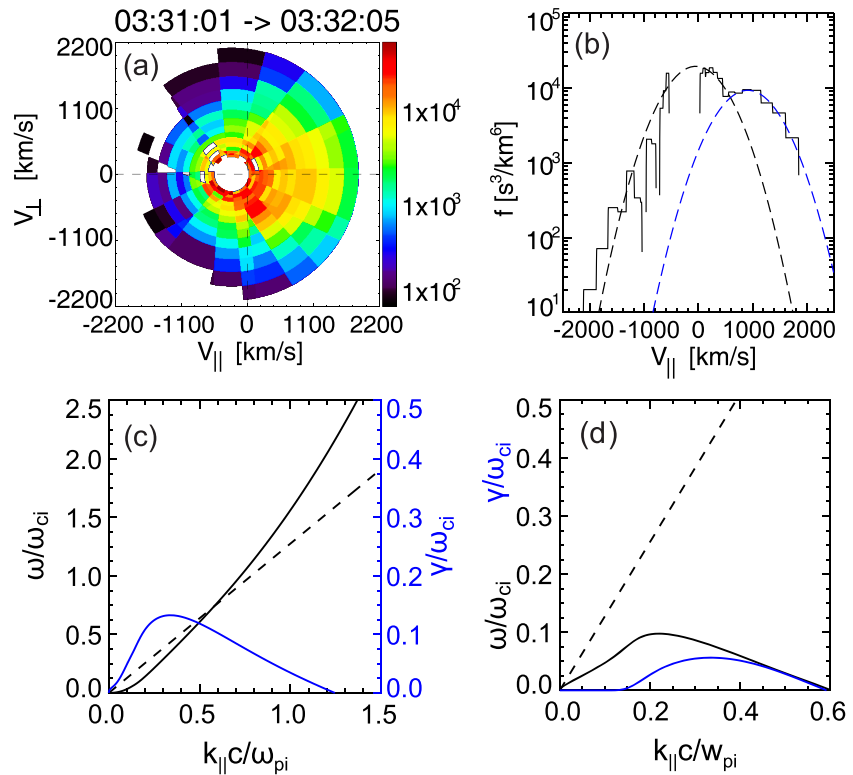




**Figure 4.** Electron distributions and wave dispersion studies of whistler mode wave 5. (a)–(e) Analysis of the electron distributions from 033156 UT to 033200 UT with the same format as in Figure 3; (f) growth rate calculated with 0.5 (blue solid line), 1 (black solid line), 2 (green solid line), 3 (red solid line) times the density and 2 (green dash line) and 3 (red dash line) times the temperature anisotropy ( $T_{\perp}/T_{\parallel}$ ) of the field-aligned populations of the electron distribution functions.

Wave dispersion analysis based on the fitted parameters shows an antiparallel propagating mode as displayed in Figure 3e. The growth rate (blue line) of this wave mode shows two positive peaks at the frequency (black line) of 0.16 and 0.23 times the electron cyclotron frequency ( $\omega_{ce}$ ). To verify the contributions to these two peaks, the four beams of the field-aligned population are removed in turn to calculate the growth rate. As shown in Figure 3f, the absence of beam 1 (blackline) will not influence the growth rate, while the disappearance of beams 2 or 3 (blue line) will result in the absence of the second peak and shift the wave number of the first peak, indicating that the second peak of the growth rate is contributed by beams 2 and 3 together. When beam 4 is removed (red line), the first peak disappeared, but the second peak became much larger, and this implies that the first peak of the growth rate is contributed by beam 4, and beam 4 could also suppress the instabilities induced by beams 2 and 3. The temperature anisotropies of the hot electrons (Figure 3d) also support a weak parallel propagating mode, but the growth rate of this mode is 30 times lower than the mode induced by the parallel enhanced population and could hardly induce waves.

The electron distribution functions of the fifth whistler mode wave have also been analyzed with the same method as shown in Figure 4. During this interval, the field-aligned population is also shown (Figure 4b), and the background electrons are isotropic (Figure 4d). The fitting results are given in the second set of Table 1, based on which wave dispersion analyses are employed. In this case, the distribution also supports an antiparallel propagating mode with a peak growth rate at the frequency of  $0.045 \omega_{ce}$  (Figure 4e). To study the influences of the parameters of the field-aligned population to the wave growth, the density and temperature anisotropy of all the beams are changed to explore their roles. As shown in Figure 4f, when the density of the field-aligned component increased from 0.5 to 3 times the initial value (shown by solid line with different colors), the growth rate and the corresponding parallel wave number  $k_{\parallel}$  were enhanced. The temperature anisotropy is changed by adjusting the perpendicular temperature of all beams (green and red dash lines). When the perpendicular temperature is half of the current distribution, no positive growth rate is shown, indicating that the excitation of this wave needs to exceed a threshold of beam temperature anisotropy. As expected, a higher temperature anisotropy with hotter perpendicular component will result in the strengthening of the growth rate and possibly the appearance of other growth rate peaks. For other whistler mode waves, the fitted parallel enhanced population of the second and fourth waves are unstable, while



**Figure 5.** Ion distribution and wave dispersion studies of the ion-scale wave. (a) Two-dimensional ion velocity distribution function from 03:31:00 UT to 03:32:05 UT in the plasma flow frame.  $V_{||}$  is the velocity parallel to the magnetic field ( $\mathbf{B}$ ), and  $V_{\perp}$  is the velocity perpendicular to  $\mathbf{B}$  and the ion bulk velocity. (b) One-dimensional cuts of ion distribution function along  $\mathbf{B}$  (black line). The black dash line shows Maxwell fitting of the background ions, while the blue dash line marks the Maxwell fitting of the ion beam. (c) and (d) are two wave modes calculated based on the ion distribution function; the frequency and growth rate are shown by black and blue solid lines, and the black dash lines mark the ion Alfvén speed ( $V_A$ ).

the distribution of the third wave did not provide positive growth rate, and this might be due to the limited time resolution of the particle instrument that could not fully record the parallel enhanced components.

The field-aligned enhanced electron populations appear periodically with the ion-scale waves (Figure 2e); hence, we suggest that they are most likely generated by the ion-scale waves. The excitation of the ion-scale wave has also been studied with particle data. The two-dimensional ion velocity distributions and one-dimensional cuts along the parallel direction in the plasma flow frame during the time of Figure 2 are shown in Figures 5a and 5b. The ion velocity distribution function shows that the phase space densities of the parallel streaming ions are much larger than that in the opposite direction. Maxwell fitting analyses show that the ion distribution consisted of a hot population (black dash line in Figure 5b) and a field-aligned ion beam (blue dash line) streaming with a velocity of  $\sim 960$  km/s, which is about 3.2 times the local Alfvén velocity ( $V_A = 297$  km/s). The fitting parameters are listed in Table 1 of last sets.

Wave analyses have been performed based on the fitting parameters, and two modes were found to be unstable. The first mode is excited by the ion beam and has a maximum growth rate when the wave vector is parallel to the magnetic field as shown in Figure 5c. The growth rate of this mode has a wide range of positive values covering the frequency from very low to  $\sim 2\omega_{ci}$  and shows a peak at the frequency of  $\sim 0.65\omega_{ci}$ . The black dash line marks the ion Alfvén speed in the  $k - \omega$  plane, and the phase velocity of this wave is close to  $V_A$  if the frequency is around  $\omega_{ci}$ . The second mode shown in Figure 5d is a low-frequency ( $\omega < 0.1\omega_{ci}$ ) wave that propagated quasiparallel to the magnetic field and generated by the temperature anisotropy of the hot ion population. Since the drift velocity of the hot ion population is low ( $0.16V_A$  according to fitting result, as shown in Figure 5b), and the flow velocity is mostly contributed by the ion beam, the Doppler shift effect of the ion-scale waves should not be large. Considering that the frequency of the ion-scale wave is close to  $\omega_{ci}$ , it

**Table 1**  
Fitting Results of the Electron and Ion Distribution Functions

033109 UT–033114 UT, $ \mathbf{B}  = 6.5$ nT, whistler mode wave				
Components	$N$ ( $\text{cm}^{-3}$ )	$V_0$ (km/s)	$T_{\parallel}$ (eV)	$T_{\perp}/T_{\parallel}$
Hot electrons	0.0435	−4,130.4	384	1.4
Cold electron	0.0258	−4,130.4	10	1.0
Electron beam 1	0.0004	5,204.6	5	6.0
Electron beam 2	0.0155	10,258.6	50	10.0
Electron beam 3	0.0068	17,166.6	60	6.0
Electron beam 4	0.0003	28,045.6	55	10.0
Ions	0.0923	0	2,000	1.0
033156 UT–033200 UT, $ \mathbf{B}  = 4.5$ nT, whistler mode wave				
Hot electrons	0.11	−2,011.3	404	1.0
Electrons beam 1	0.012	16,532.7	59	8.0
Electrons beam 2	0.00075	25,817.7	41	8.0
Electrons beam 3	0.00010	34,862.7	64	8.0
Ions	0.12285	0	2,000	1.0
03:31:00 UT–03:32:05 UT, $ \mathbf{B}  = 4.9$ nT, ion-scale wave				
Hot Ion	0.047	−189	2,219	1.56
Cold Ion	0.059	−189	100	1.0
Ion Beam	0.026	780	2,346	1.66
Electrons	0.132	0	487	1.0

Note.  $N$ ,  $V_0$ ,  $T_{\parallel}$ , and  $T_{\perp}/T_{\parallel}$  indicate density, beam streaming velocity, parallel temperature, and temperature anisotropy.

is more likely to be induced by the first mode with the higher frequency generated by the ion beams. In observations, the fluctuations of the ion-scale waves were mainly in the perpendicular direction (Figure 2a), and this resulted in a quasiparallel propagation mode that is consistent with the wave analyses.

#### 4. Discussion and Summary

Taking account of the properties presented above, a possible candidate of the ion-scale waves is the kinetic Alfvén wave (KAW) (Hasegawa & Chen, 1976). KAW can be generated by ion–ion streaming (Winske & Omid, 1992; Yin et al., 2007) and carry a considerable amount of energy away from the fast plasma sheet flows (Chaston et al., 2012). KAW occurs when the perpendicular scale of the Alfvén wave (Alfvén, 1942) approaches the ion cyclotron scale, and the frequency of KAW is of the order of  $f_{ci}$ . One important feature of the KAW is that it can generate parallel electric field, and we interpret that the enhanced parallel electron phase space density is a result of trapping and acceleration of the electrons by the parallel electric field (Artemyev et al., 2015; Hui & Seyler, 1992; Kletzing, 1994). However, the electric field data during the wave interval are complicated and could have included noise, making it difficult to establish the properties of the KAW with certainty. Therefore, we leave this issue for future studies.

In this event, the enhancement of the field-aligned component only appears in the parallel direction, and no counterstreaming beams are observed. We interpret this character as an effect of the magnetic field configuration on the tailward side behind antidipolarization front (at 033000UT in Figure 1a), where the magnetic field is connected to the interplanetary magnetic field, and the field-aligned populations will not bounce back.

In the Earth’s magnetosphere, whistler mode waves can be generated by electron temperature anisotropy ( $T_{\perp} > T_{\parallel}$ ) (Kennel & Petschek, 1966; Li et al., 2018) or electron beams (Ren et al., 2019; Sauer & Sydora, 2010). In this study, the whistler mode waves are generated by electron populations streaming on one side along the magnetic field and such a mechanism resembles the whistler heat flux instability (WHFI) (Kuzichev et al., 2019; Tong, Vasko, Marc, et al., 2019) observed in the solar wind. The electron beta  $\beta_e$  (electron thermal pressure vs. the magnetic pressure) in this event varied from 1 to 6 (Figure 1h), within the statistical range of the instability observed in the solar wind (Tong, Vasko, Artemyev, et al., 2019), and the properties of these whistler waves that include the frequency range, right-hand circular polarization, and quasiparallel propagation are all similar to that generated by the



solar wind WHFI (Tong, Vasko, Marc, et al., 2019). All these properties support the similarity of generation mechanism of the whistler mode waves discussed in this paper to the WHFI. As far as we know, the heat flux instability has not been reported in the Earth's magnetosphere. Hence, the observations in this study give information on the heat flux instability exhibited in the distant magnetotail but with the help of ion-scale waves.

Zhao et al. (2017) have shown that modulation of whistler mode waves also occurs near the neutral sheet region of distant magnetotail. However, the wave excitation and modulation mechanisms are different from the observations presented in this article. Near the neutral sheet region, the whistler mode waves are generated by temperature anisotropy of the electron flat-top distributions, and the ion-scale waves could confine the generation and propagation of the whistler mode waves near the magnetic field valley. In this study, similar processes were also observed from 03:30:10 UT to 03:30:30 UT when the spacecraft crossed the neutral sheet region (Figure 1). Since the major conclusions are similar, the detail analysis was omitted here. These observations imply that the modulation of whistler mode waves by ion-scale waves can take place with different mechanisms in the plasma sheet.

To reiterate, we have observed wave modulations between whistler mode waves and ion-scale waves in the plasma sheet at  $\sim 54 R_E$ . The ion-scale wave, most likely KAW, are excited by the interaction of the ion beam and the background ions. The ion-scale waves periodically generated field-aligned electron populations, which drove whistler mode waves at each cycle of the ion-scale waves. These observations have suggested that plasma dynamics on electron and ion-scales can be coupled through waves.

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